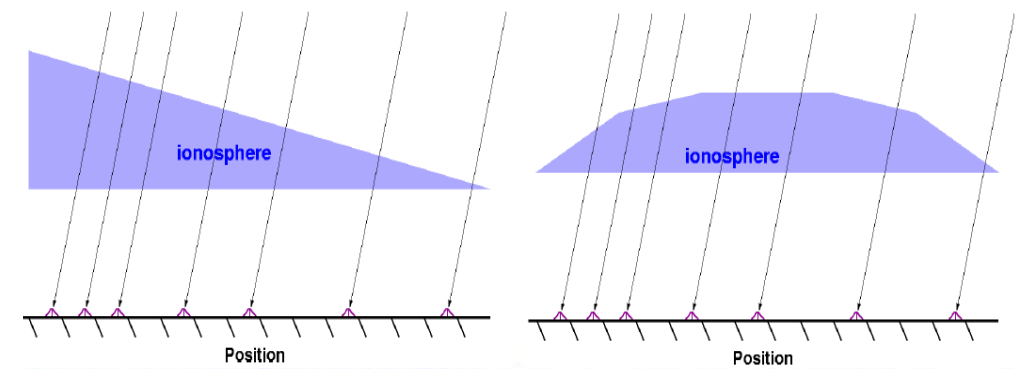
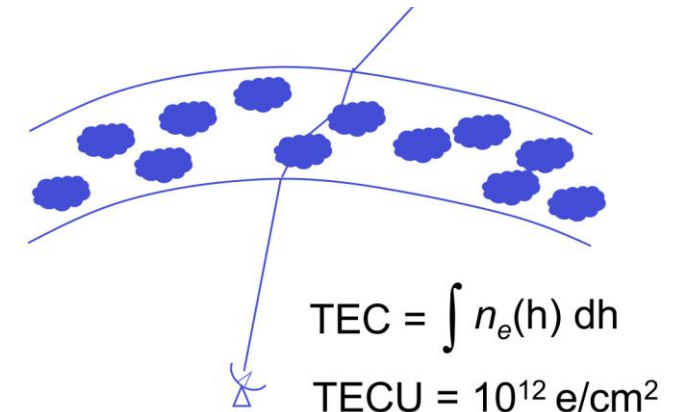


Fitting and Testing Ionospheric Phase Screens with MSSS Data

David Rafferty, Bas van der Tol, George Heald (+TT)

Introduction

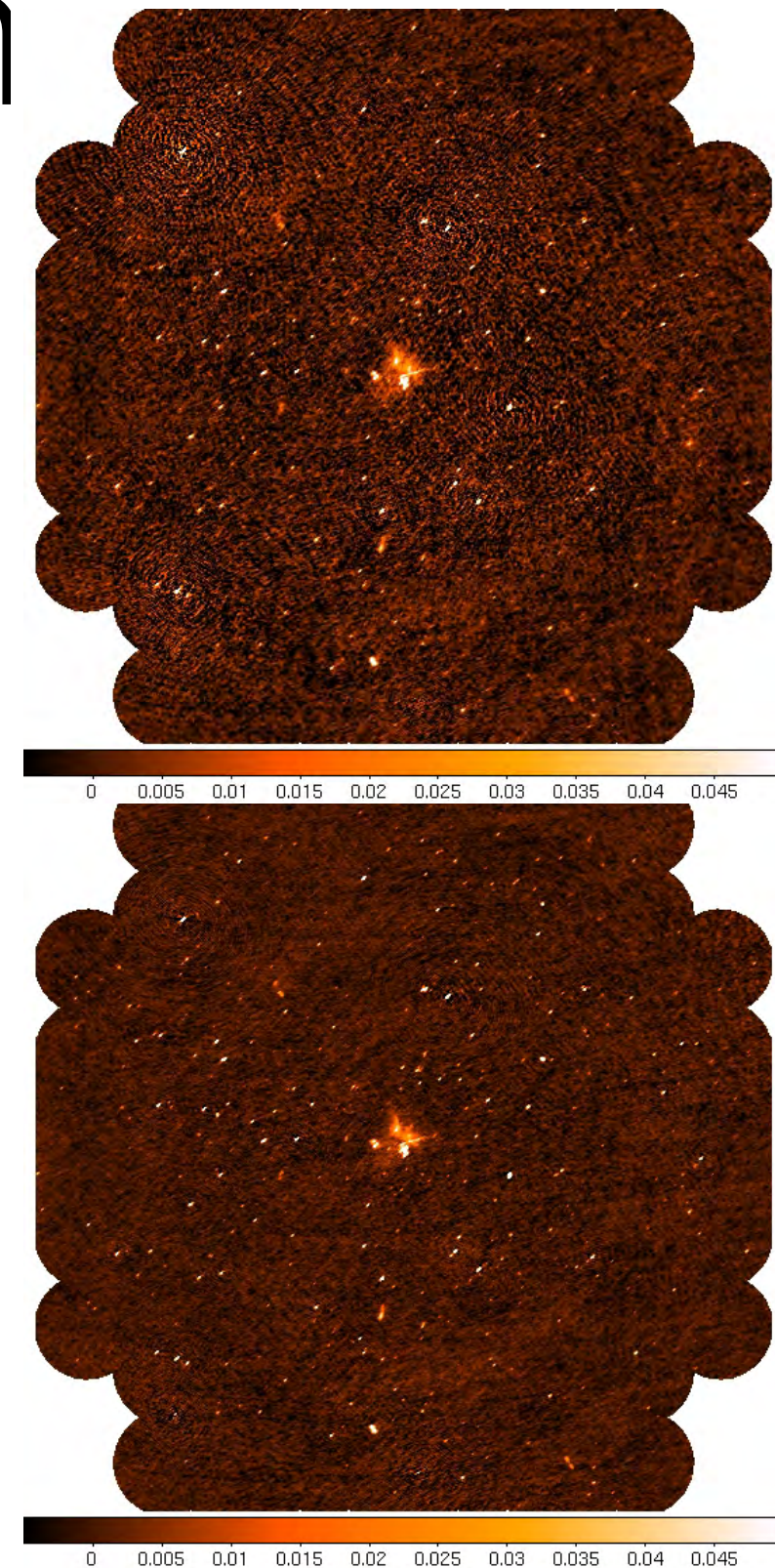
- The ionosphere can cause time- and position-dependent phase shifts
- Using direction-dependent calibration, these shifts can be measured and corrected (e.g., with phase screens as in SPAM)
- The SPAM approach assumes that instrumental effects have been removed, but (so far) this has not been possible with LOFAR
- One solution: use phase differences between sources



Credit: Huib Intema

Introduction

- The ionosphere can cause time- and position-dependent phase shifts
- Using direction-dependent calibration, these shifts can be measured and corrected (e.g., with phase screens as in SPAM)
- The SPAM approach assumes that instrumental effects have been removed, but (so far) this has not been possible with LOFAR
- One solution: use phase differences between sources

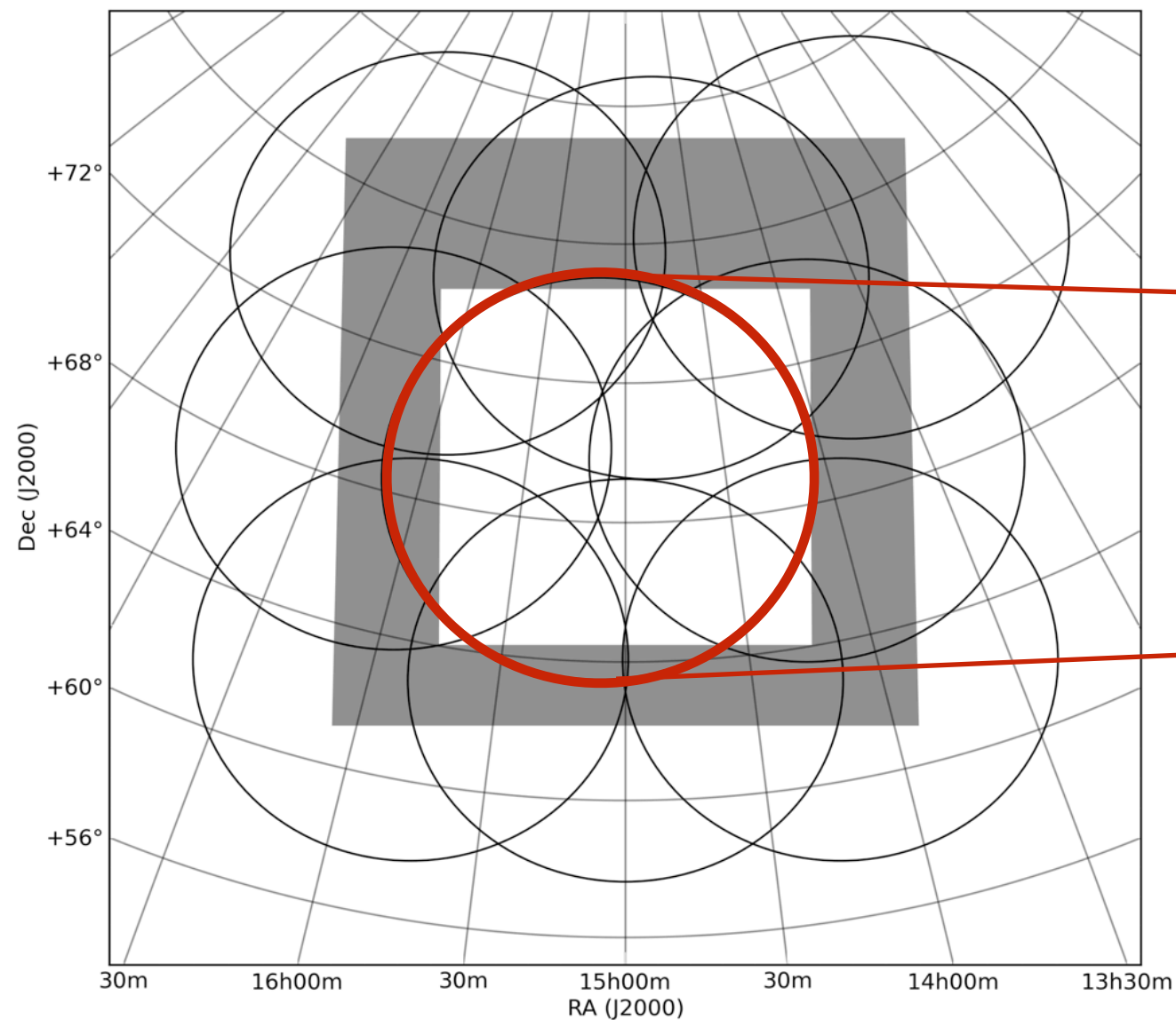


Credit: Huib Intema

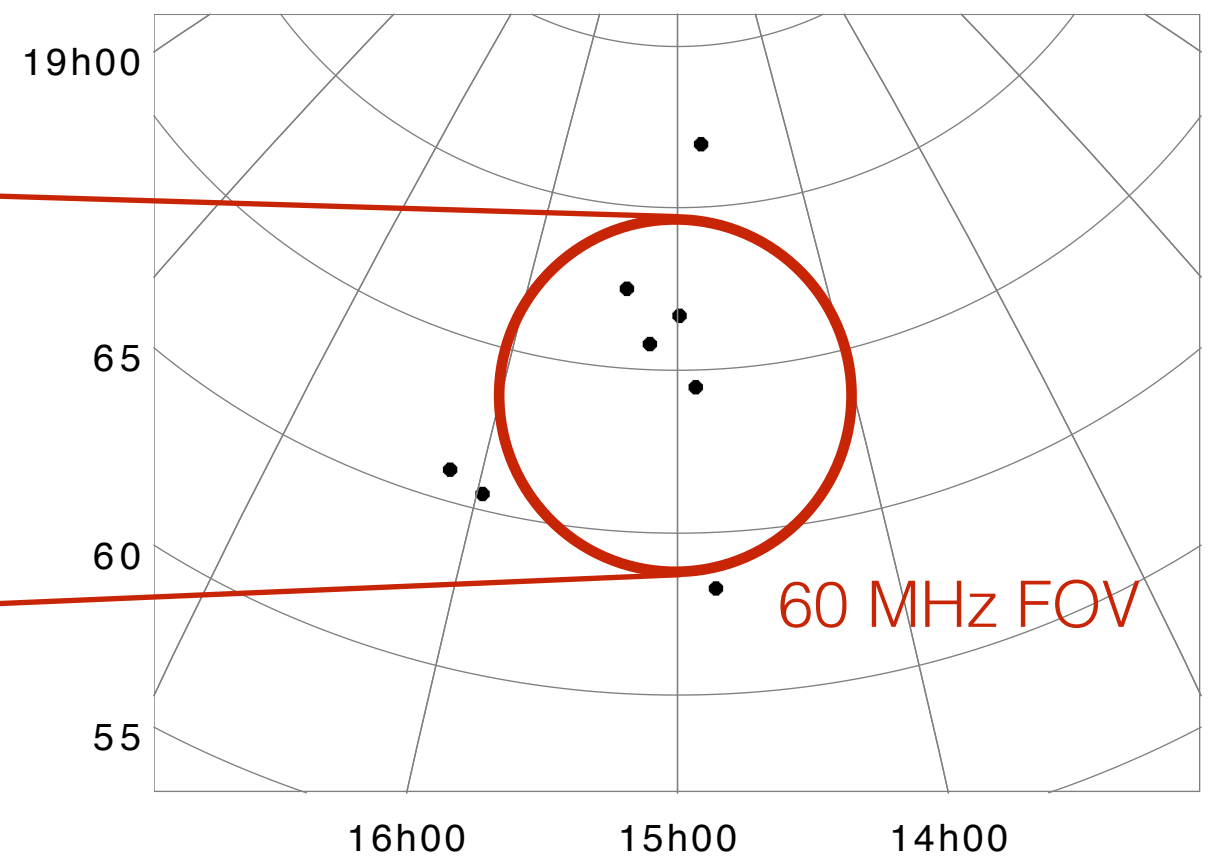
Source Differencing

- Perform direction-dependent calibration for bright sources
- Assume that instrumental effects are the same in all directions
- Subtracting phase solutions for two sources will result in purely direction-dependent (ionospheric) effects
- Test with MSSS (MVF) LBA data: 8 2-MHz bands, 9 11-minute snapshots

MSSS Verification Field

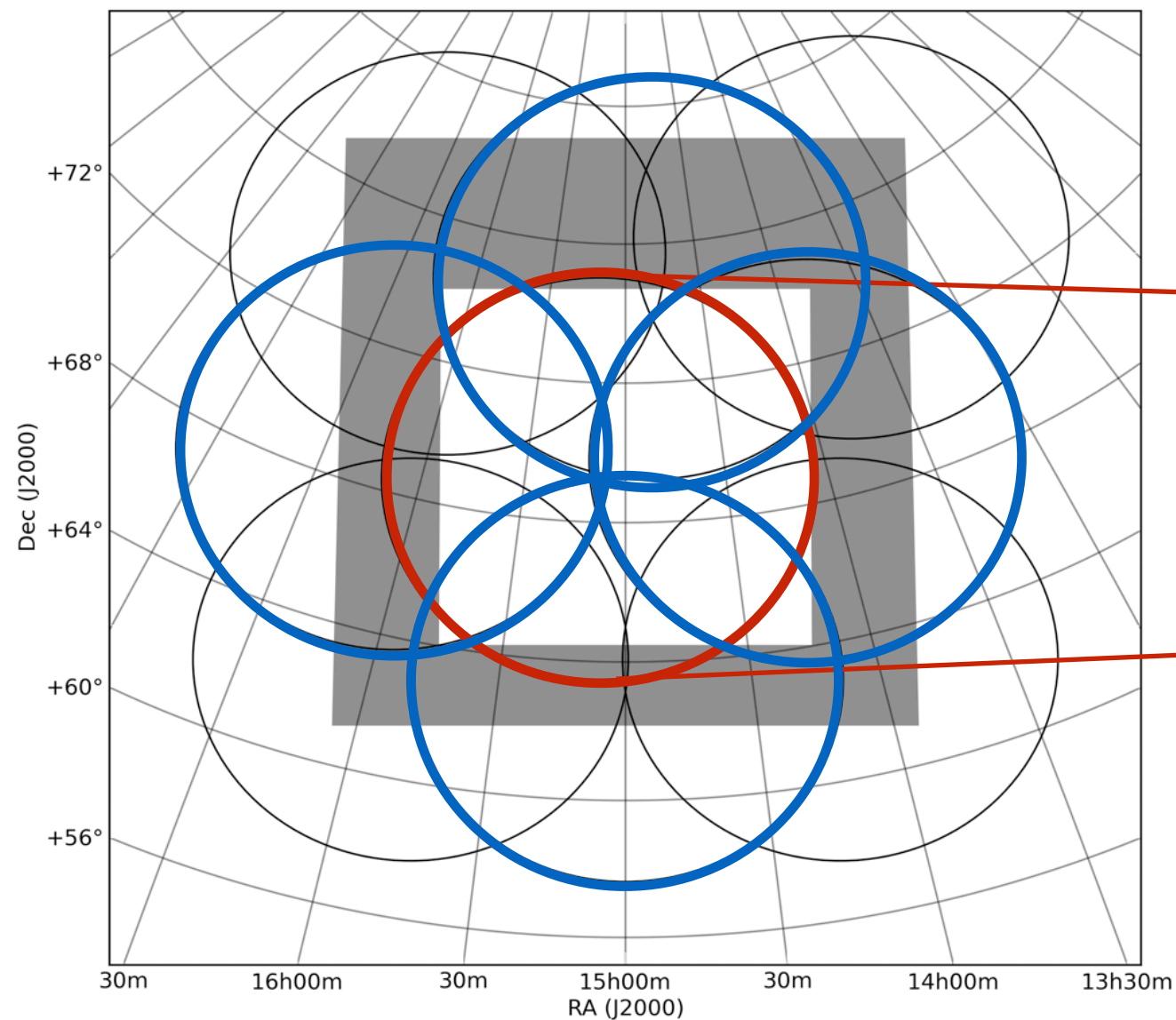


MSSS Verification Field layout

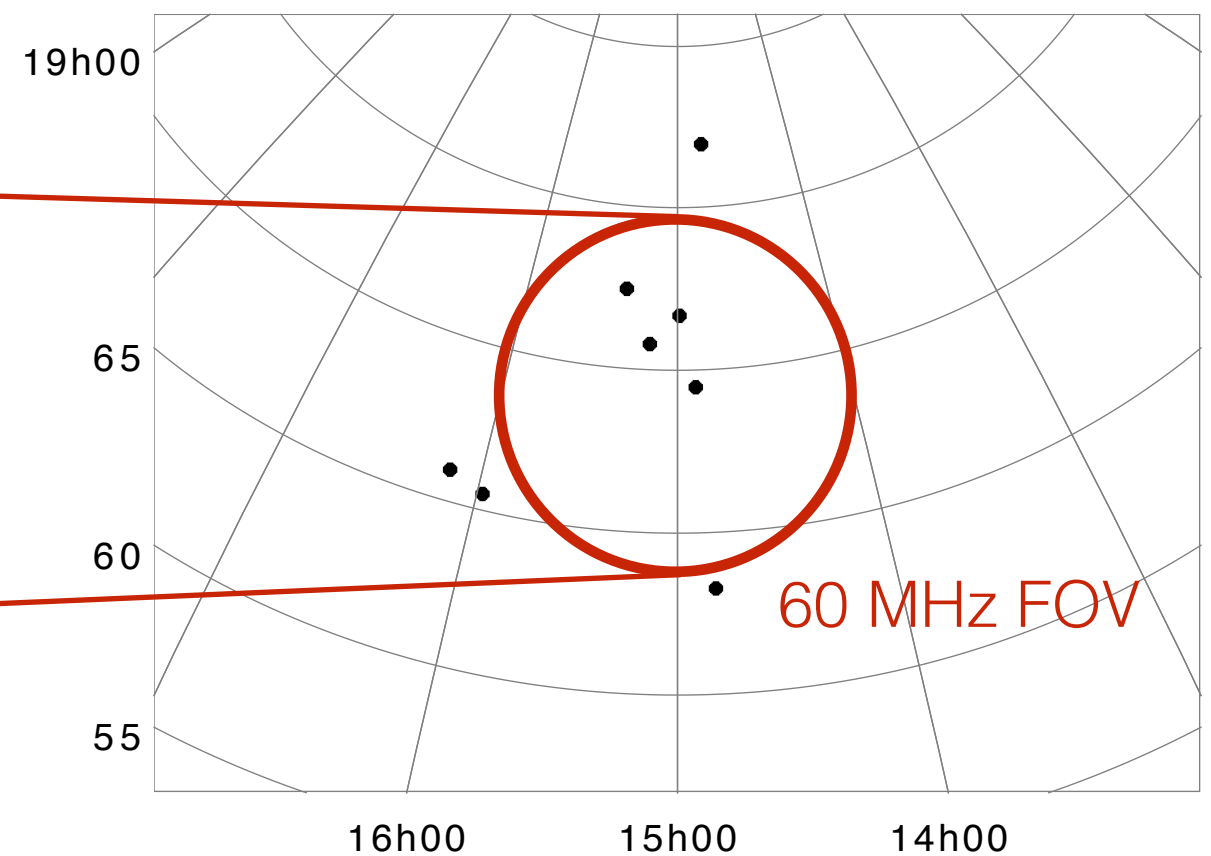


Phase-screen calibrators
($S_{60} > 10$ Jy)

MSSS Verification Field



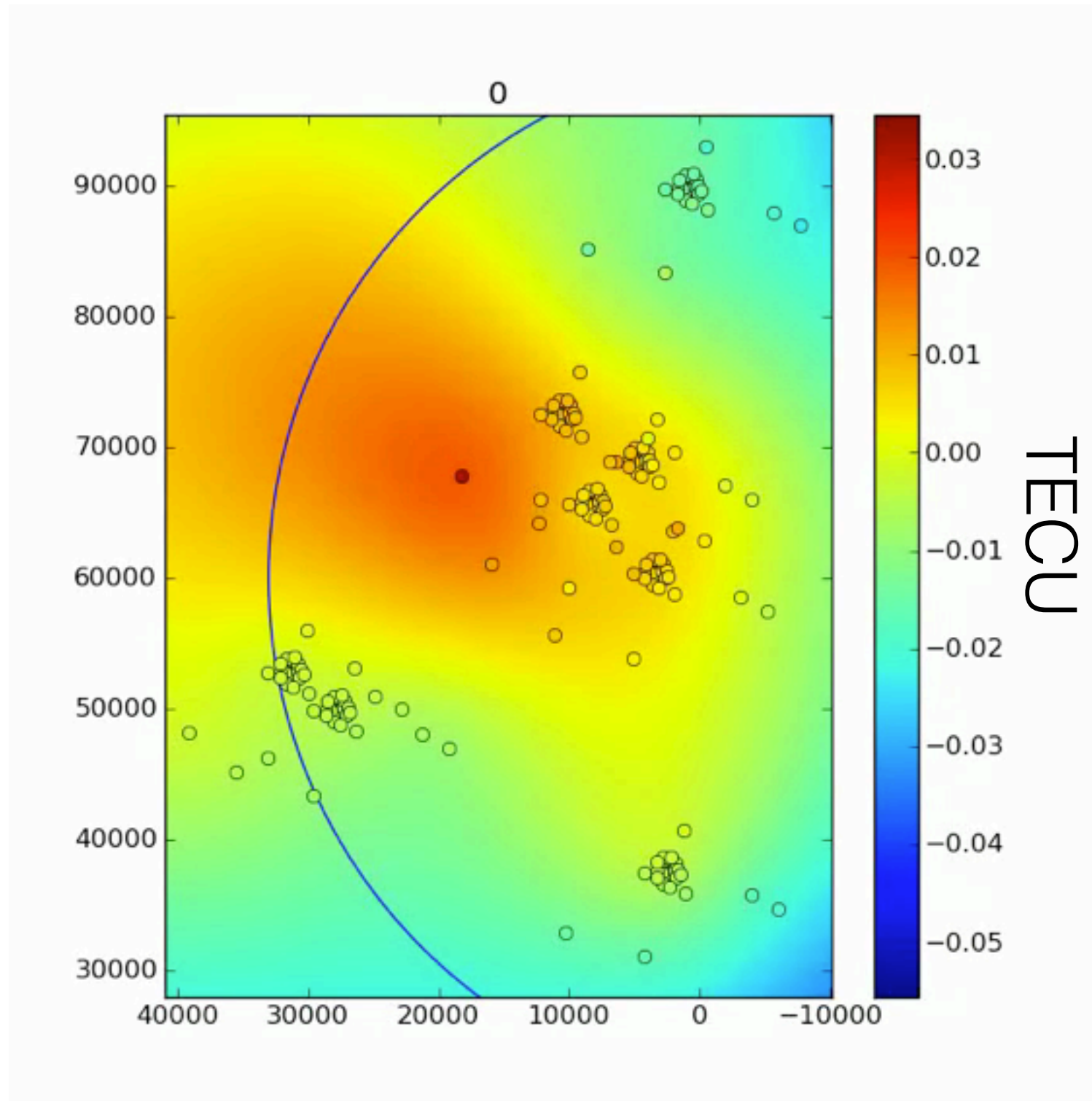
MSSS Verification Field layout



Phase-screen calibrators
($S_{60} > 10$ Jy)

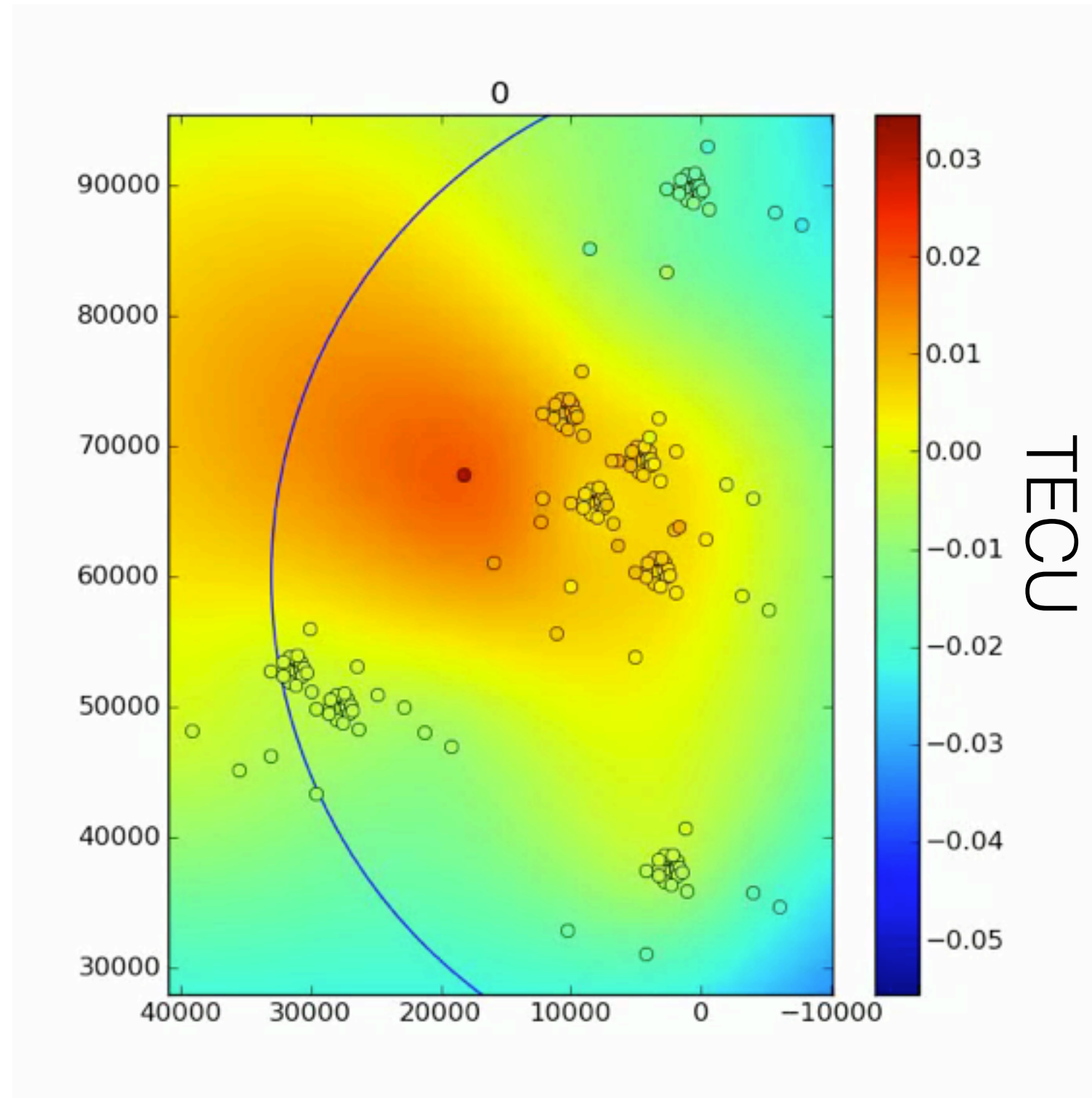
Example Phase Screen

- TEC value was derived for each pierce point every 10 seconds using fit to phases across all 8 bands
- All core stations + 5 remote stations were used
- 7 11-minute snapshots were used (first two snapshots not used due to poor solutions)

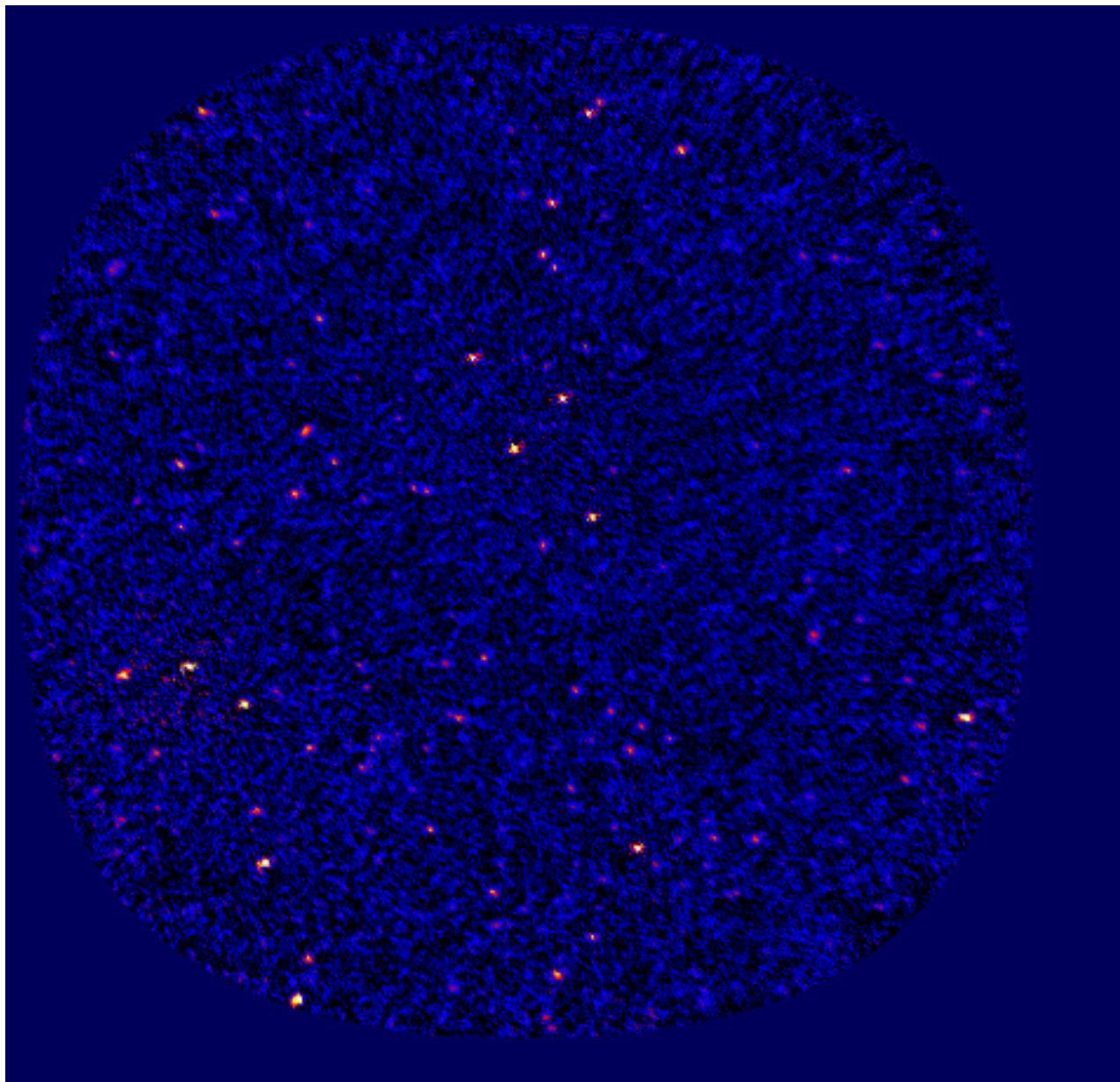


Example Phase Screen

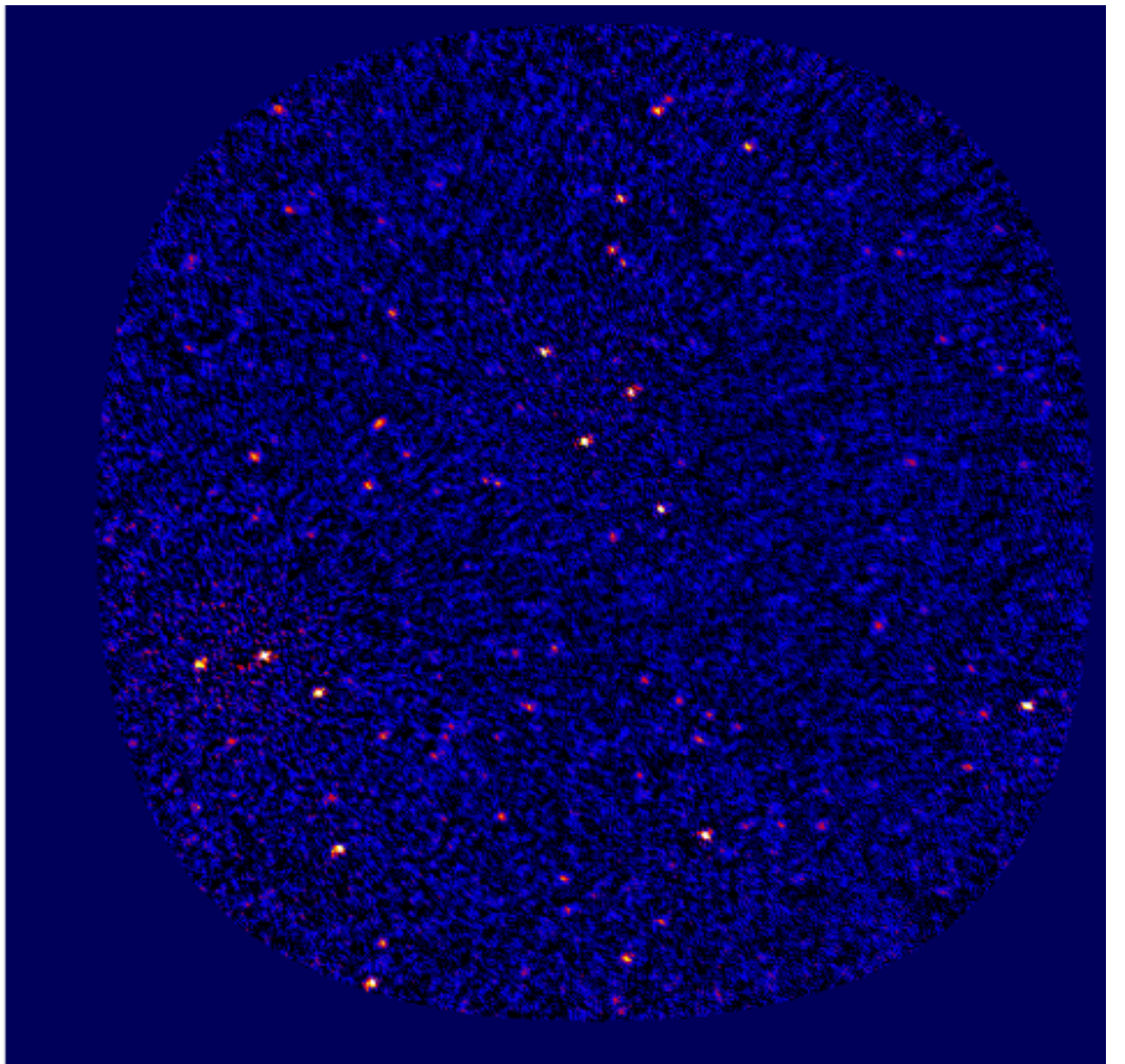
- TEC value was derived for each pierce point every 10 seconds using fit to phases across all 8 bands
- All core stations + 5 remote stations were used
- 7 11-minute snapshots were used (first two snapshots not used due to poor solutions)



30 MHz Images

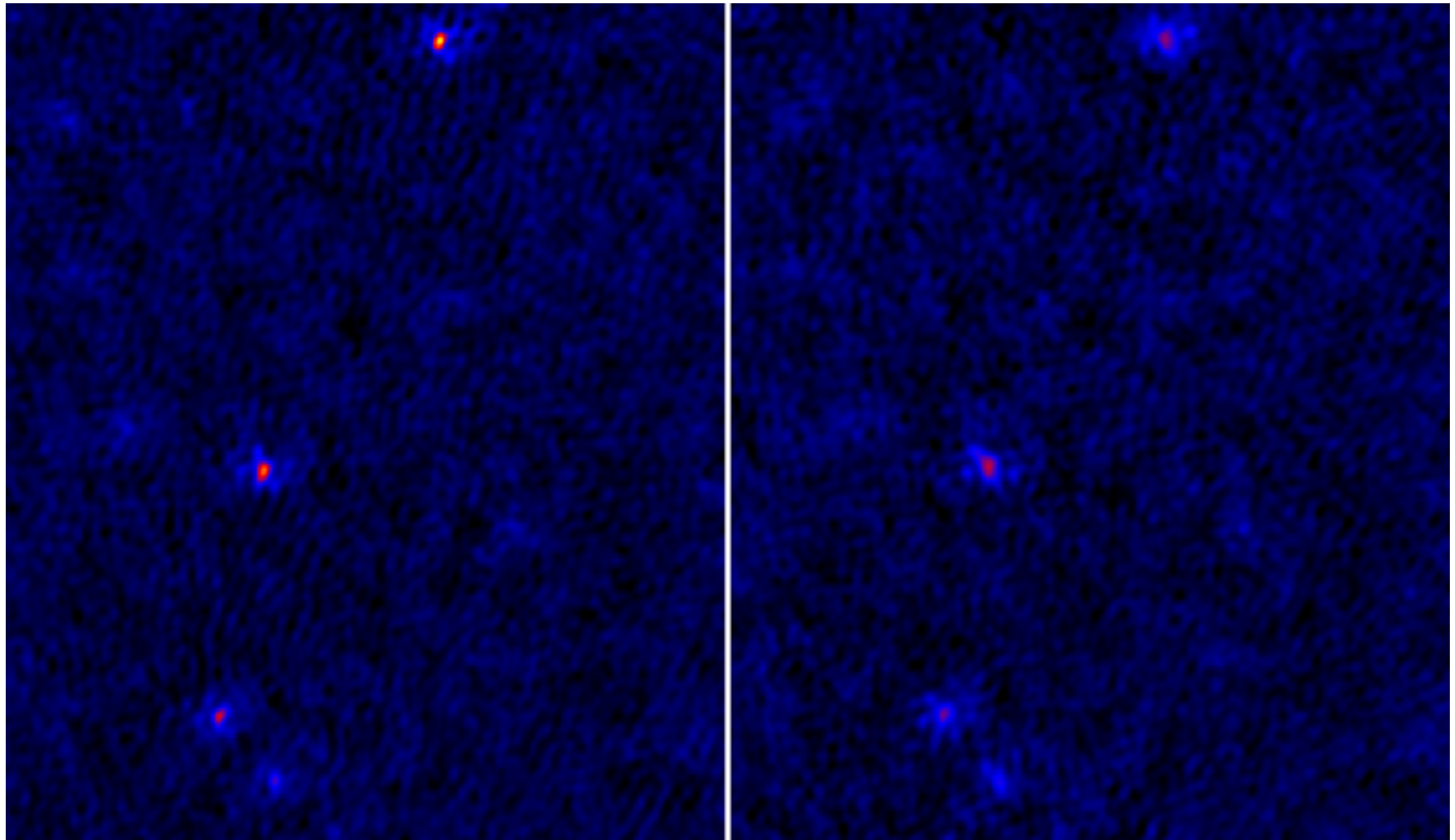


With phase screen



Without phase screen

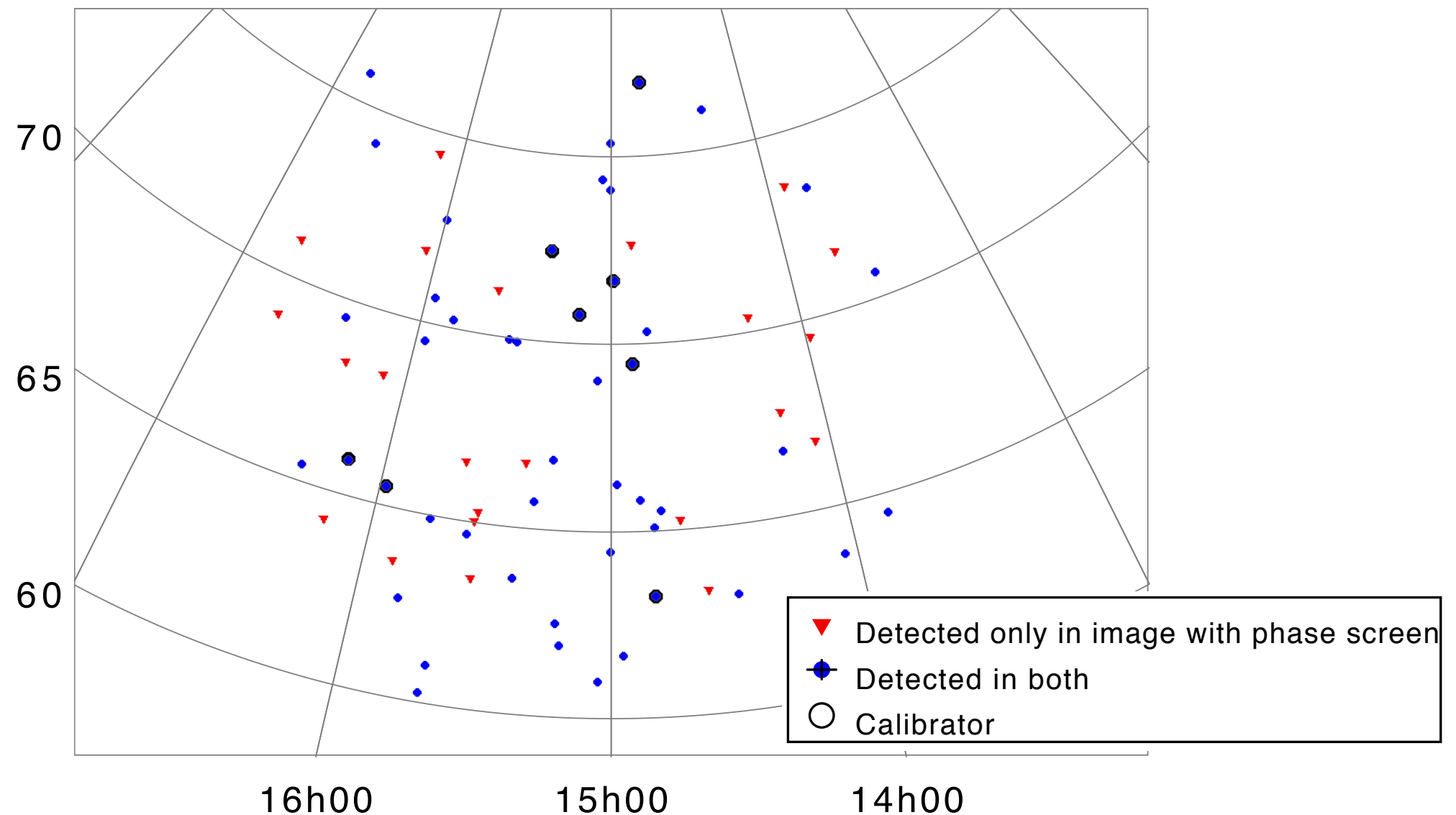
30 MHz Images



With phase screen

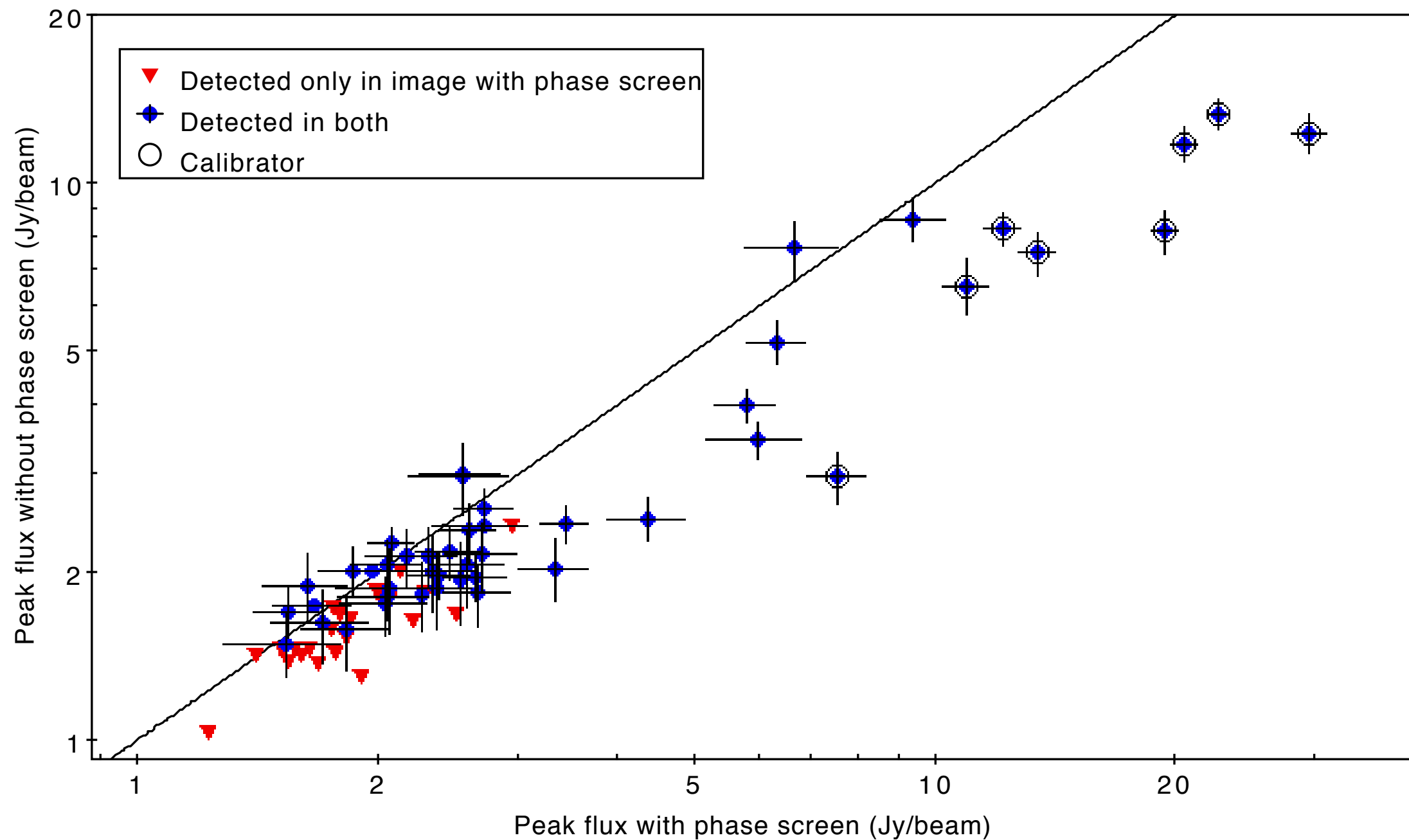
Without phase screen

Detected Sources at 30 MHz ($>6\sigma$ peak flux)

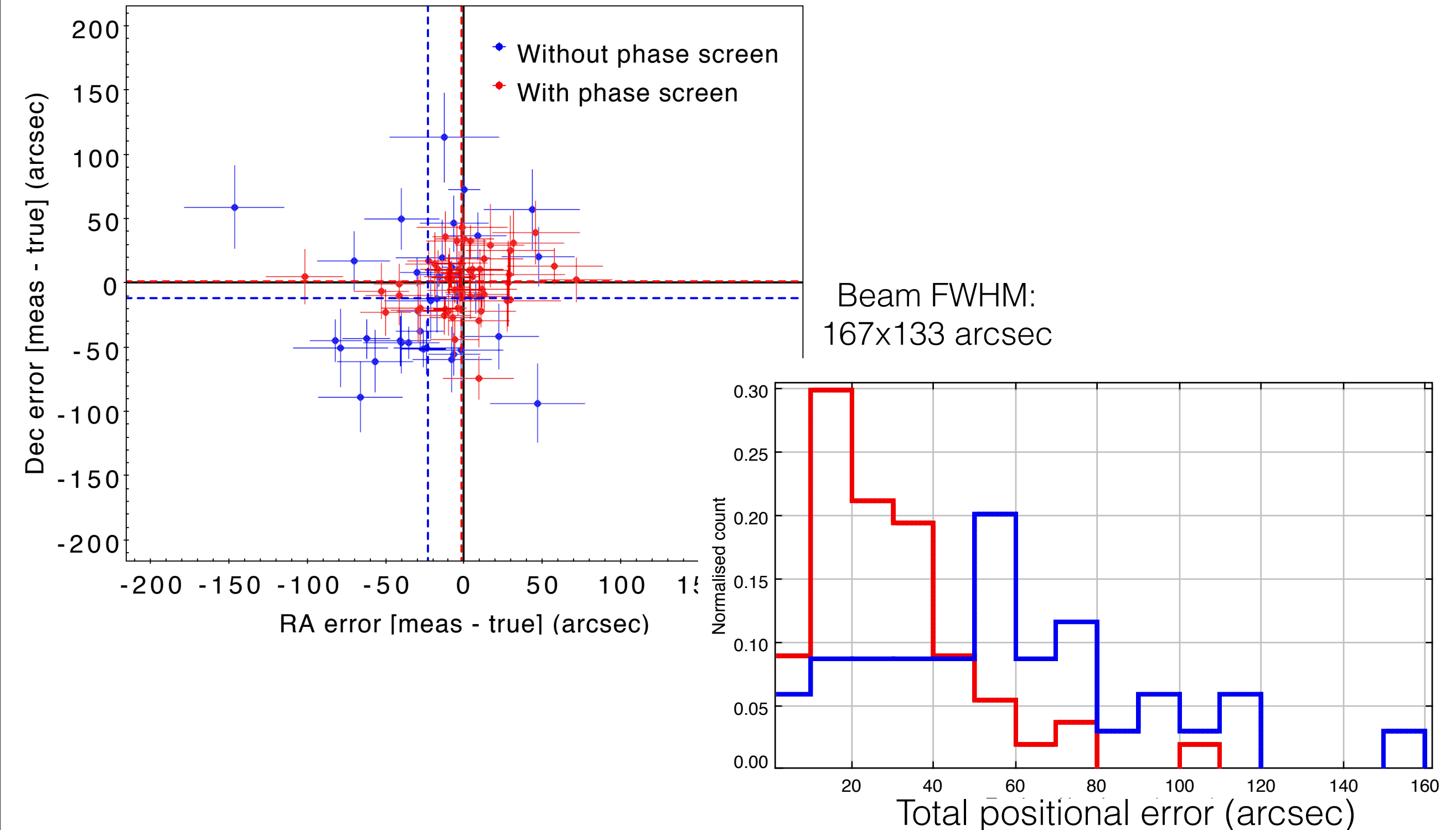


- At 30 MHz, $\sim 50\%$ more sources detected in image with phase screen ($\sim 30\%$ more at 45 MHz)

Peak Fluxes at 30 MHz



Positional Errors at 30 MHz



To-do and Application to Other LBA Observations

- Image all 8 bands and compare source detection to images without screen (Giulia)
- Try more sophisticated peeling strategies (varying solution intervals, use of patches, etc.)
- Investigate different screen heights and two layers
- Likely need simultaneous flanking field observations to obtain enough calibrators in all bands, so considerable bandwidth may be required
- Can require a lot of time: current approach for 8 bands, 11 minutes = 10-100 hours